The interactions of winds from massive young stellar objects

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Abstract: The supersonic stellar and disk winds possessed by massive young stellar objects will produce shocks when they collide against the interior of a pre-existing bipolar cavity (resulting from an earlier phase of jet activity). The shock heated gas emits thermal X-rays which may be observable by spaceborne observatories such as the Chandra X-ray Observatory. Hydrodynamical models are used to explore the wind-cavity interaction. Radiative transfer calculations are performed on the simulation output to produce synthetic X-ray observations, allowing constraints to be placed on model parameters through comparisons with observations. The model reveals an intricate interplay between the inflowing and outflowing material and is successful in reproducing the observed X-ray count rates from massive young stellar objects.

1 Introduction

Observational and theoretical advances have provided increasing evidence that massive star formation is not merely a scaled-up version of lower-mass star formation. However, there are some similarities. For instance, both involve outflows (Garay & Lizano 1999; Reipurth & Bally 2001; Banerjee & Pudritz 2006, 2007) and bipolar cavities are commonly observed around both high and low mass young stellar objects (Garay & Lizano 1999). The widths of IR recombination line emission observed from massive young stellar objects (MYSOs) indicates the presence of dense ionized outflows with velocities ranging from 100 to > 340 km s^{-1} (Drew, Bunn, & Hoare 1993; Bunn, Hoare, & Drew 1995). Further confirmation of outflows has come from high angular resolution radio observations (e.g. Hoare et al. 1994; Hoare 2006; Curiel et al. 2006). One explanation would be a disk wind driven by radiation pressure at the surface of the disk (e.g. Drew et al. 1998).

X-rays have been detected from deeply embedded MYSOs in star forming regions by the *Chandra X-ray observatory* (hereafter *Chandra*; e.g. Broos et al. 2007; Wang et al. 2007). Preibisch et al. (2002) found that the X-ray emission from IRS3 A and C, with a count rate of 0.166 ± 0.041 ks^{-1} for the former, could not be explained by the standard scenario for massive stars (i.e. wind embedded shocks produced by instabilities inherent in radiatively-driven winds - see Owocki, Castor, & Rybicki 1988), yet the estimated stellar masses of these objects implies they will have radiative outer envelopes which poses problems for the generation of X-rays through magnetic star/disk interactions.

A potential source of X-rays which has not been considered before is the collision between the stellar and disk winds and the infalling envelope. We explore this scenerio using hydrodynamical models and find that the interaction of the stellar and disk wind with the cavity wall produces shock heated plasma which emits X-rays in agreement with Chandra observations (Parkin et al. 2009).

The remainder of this work is structured as follows: in Sect. 2 we describe the winds-cavity model, in Sect. 3 we present the result of the hydrodynamic simulations and X-ray calculations, and in Sect. 4 we close with our conclusions.
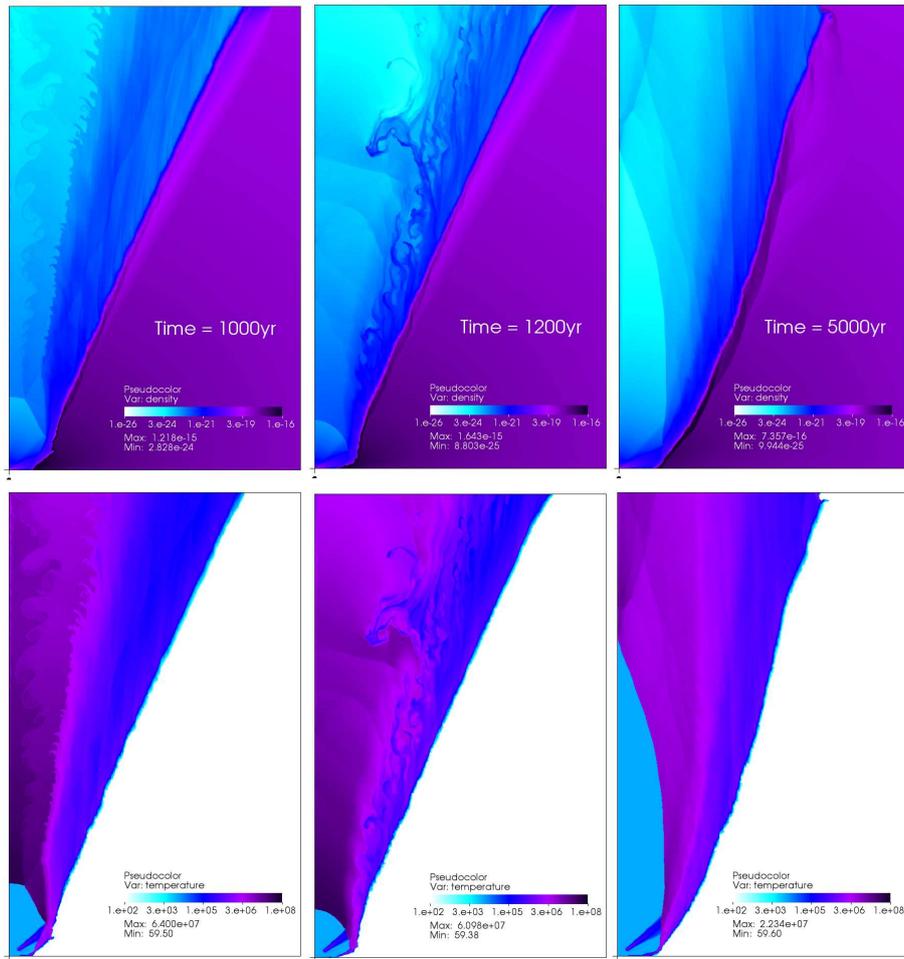

Figure 1: Density (top row) and temperature snapshots (bottom row) from our fiducial model at a simulation time of $t = 1000$ (left), 1200 (middle), and 5000 yrs (right). The grids extend to $x = 5 \times 10^{16}$ cm and $y = 8 \times 10^{16}$ cm. The central star is located in the bottom left corner. The reverse shock can be seen in the lower left of the grid, where the enclosed preshock wind has a temperature of $T = 10^4$ K.

2   The winds-cavity interaction model

The model consists of a MYSO situated at the centre of a previously evacuated bipolar cavity which is surrounded by infalling molecular cloud material. We include the stellar wind and a disk wind which emanates from the surface of the circumstellar accretion disk; both are assumed to be at terminal velocity. Due to the spatial scales under consideration we do not attempt to model the radiative driving of the winds as this requires high spatial resolution in the vicinity of the star/disk (e.g. Proga

et al. 1998). For simplicity we adopt an angle dependent wind prescription based on the models of Proga et al. (1998), Drew et al. (1998), and Sim et al. (2005), whereby the stellar wind occupies a region from a polar angle $0 - 60°$ and the disk wind occupies the region from $60° - 85°$. The density and velocity distibutions for the infalling cloud material are described by the equations of Terebey et al. (1984), and a simple analytical prescription similar to that of Alvarez et al. (2004) is used to determine the morphology of the pre-existing outflow cavity. We model the winds-cavity interaction in 2D cylindrical symmetric coordinate system using VH-1 (Blondin et al. 1990). The code uses the piecewise-parabolic method of Colella & Woodward (1984) to solve the hydrodynamic equations on a fixed grid. Further details of the model can be found in Parkin et al. (2009).

We consider a 30 $M_{\odot}$ O8V star with a mass-loss rate of 10^{-7} $M_{\odot}$ yr^{-1} and terminal wind speed of 2000 km s^{-1}. The disk wind has a mass-loss rate of 10^{-6} $M_{\odot}$ yr^{-1} and a terminal wind speed of 400 km s^{-1}. The unshocked winds are assumed to be at a temperature of 10^4 K. The mass infall rate for the cloud is 2×10^{-4} $M_{\odot}$ yr^{-1} and the cavity opening angle is 30° . The winds-cavity interaction was followed for a simulation time of $t = 5000$ yrs.

To allow a comparison to be made between *Chandra* X-ray observations of MYSOs and the simulations we calculate attenuated X-ray fluxes. The emissivity values used are for optically thin gas in collisional ionization equilibrium obtained from look-up tables calculated from the MEKAL plasma code (Liedahl, Osterheld, & Goldstein 1995 and references there-in). Ray-tracing calculations are performed with an inclination angle (to the pole) of 60°.

In this work we only consider one set of parameters, and we refer the reader to Parkin et al. (2009) for a detailed parameter space exploration.

3   Results

Fig. 1 shows the spatial distribution of material in the simulation; the disk wind lines the cavity wall and separates the stellar wind from the infalling molecular cloud material. The winds in the simulation are supersonic, and their collision against the cavity wall generates a reverse shock close to the star (< 500 AU). Because the ram pressure of the inflow/outflow is angle dependent and the base of the cavity is subject to instabilities the position of the reverse shock oscillates, and its shape is often non-spherical. This is due to small fluctuations in the shape and size of the base of the cavity wall as inflowing material is ablated and incorporated into the outflow, and as new inflowing material replenishes it (see Fig. 1). The shear layer between the stellar and disk winds provides a site for the growth of ~ 700 AU (10^{16} cm) amplitude Kelvin-Helmholtz instabilities on timescales of $\sim$ a few years. By $t = 1200$ yrs an instability of this proportion can be seen driving a clump of disk wind material into the path of the stellar wind, which leads to mass-loading of the latter.

At a simulation time of $t = 1000$ yrs, the intrinsic X-ray emission comes mainly from disturbed cloud material ($L_{\text{int}_\text{C}} \simeq 3.3 \times 10^{33}$ erg s^{-1}), then the disk wind ($L_{\text{int}_\text{D}} \simeq 4.3 \times 10^{32}$ erg s^{-1}), and finally the stellar wind ($L_{\text{int}_\text{S}} \simeq 5 \times 10^{30}$ erg s^{-1}). The shock driven into the cloud is too slow to heat gas up to X-ray emitting temperatures, but large quantities of cloud material are ablated by the outflowing disk wind and mixed into this hotter flow[1]. This process heats the cloud material to temperatures where (soft) X-rays are emitted. However, the difference between the *attenuated* luminosities of the cloud and stellar wind components is very small ($L_{\text{att}_\text{C}} \simeq 8 \times 10^{28}$ erg s^{-1} and $L_{\text{att}_\text{S}} \simeq 7 \times 10^{28}$ erg s^{-1}, respectively). The explanation is that the stellar wind emission is harder and extends to higher energies, and is less affected by attenuation. In contrast, the cloud material, which is heated to lower temperatures, emits prolifically at low energies and has a much higher intrinsic

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[1]Note that some heating of the cloud material may occur due to numerical heat conduction between the hot postshock disk wind and the cooler cloud material (Parkin & Pittard 2010).
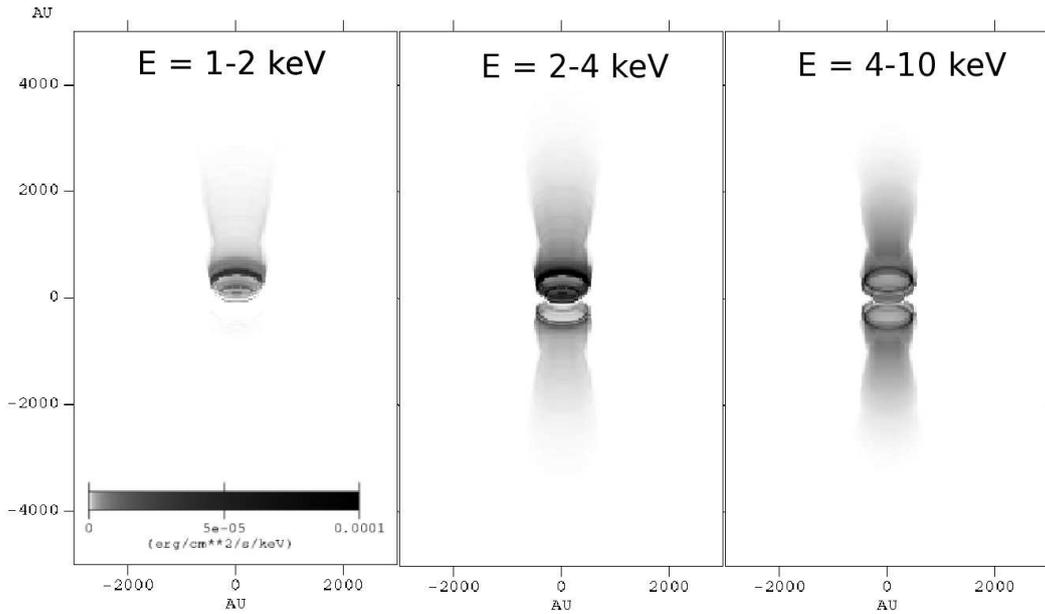
Figure 2: Ray-traced synthetic broadband X-ray images.

luminosity, but its emission is subject to considerable attenuation at $E < 1$ keV. Although some variation in X-ray emission occurs due to fluctuations in the position of the reverse shock, these values are indicative of the mean luminosities from the simulation.

The total attenuated luminosity from the model is $L_{\mathrm{att_{tot}}} \simeq 2 \times 10^{29}$ erg s^{-1} which, when placed at a distance of 1 kpc and convolved with the *Chandra* effective area, equates to a count rate of $\simeq$ 0.1 ks^{-1}. Interestingly, this is in approximate agreement with X-ray count rates of 0.166 ± 0.041 ks^{-1} and 0.30 ± 0.11 ks^{-1} inferred from observations of Mon R2 IRS 3A by Preibisch et al. (2002) and for S106 IRS4 by Giardino, Favata, & Micela (2004), respectively.

Examining synthetic broadband X-ray images calculated from the simulation output one can see that the observable X-ray emission in the 1-2, 2-4, and 4-10 keV bands originates from similar regions of the cavity (Fig. 2). The peak in the intensity in the three bands have common origins near the reverse shock, and are mainly generated by shocked stellar and disk wind material. Importantly, the spatial extents of the detectable emission in all three energy bands (1-2, 2-4, and 4-10 keV) are just below the resolution limit of *Chandra* ($\sim 0.5''$) and so this model is consistent with the unresolved nature of real MYSOs.

4   Conclusions

The wind-cavity interaction around an embedded MYSO has been studied using hydrodynamic simulations and X-ray calculations. In summary, the collision of the winds against the cavity wall generates a reverse shock close to the star (< 500 AU). The shock heated gas produces X-ray emission with an integrated count rate ($\simeq 0.1$ ks^{-1}) and spatial extent ($< 0.5''$) in agreement with observations of MYSOs by *Chandra*.

We close with a note that future X-ray satellites, such as the International X-ray Observatory (IXO), may have the potential to resolve the spatial extent of the X-ray emission from the winds-cavity interaction. If this were the case, the winds-cavity model could be used in concert with detailed analysis at infra-red wavelengths to place unsurpassed constraints on the parameters of MYSO outflows.

Acknowledgements

References

Alvarez, C., Hoare, M. & Lucas, P. 2004, A&A, 419, 203

Banerjee, R. & Pudritz, R. E. 2006, ApJ, 641, 949

Banerjee, R. & Pudritz, R. E. 2007, ApJ, 660, 479

Blondin, J. M., Kallman, T. R., Fryxell, B. A. & Taam, R. E. 1990, ApJ, 356, 591

Broos, P. S., Feigelson, E. D., Townsley, L. K., Getman, K. V., Wang, J., Garmire, G. P., Jiang, Z. & Tsuboi, Y. 2007, ApJS, 169, 353

Bunn, J. C., Hoare, M. G. & Drew, J. E. 1995, MNRAS, 272, 346

Colella, P. & Woodward, P. R. 1984, J. Comput. Phys., 54, 174

Curiel, S., Ho, P. T. P., Patel, N. A., Torrelles, J. M., Rodríguez, L. F., Trinidad, M. A., Cantó, J., Hernández, L., et al. 2006, ApJ, 638, 878

Drew, J. E., Bunn, J. C. & Hoare, M. G. 1993, MNRAS, 265, 12

Drew, J. E., Proga, D. & Stone, J. M. 1998, MNRAS, 296, L6

Garay, G. & Lizano, S. 1999, PASP, 111, 1049

Giardino, G., Favata, F. & Micela, G. 2004, A&A, 424, 965

Hoare, M. G. 2006, ApJ, 649, 856

Hoare, M. G., Drew, J. E., Muxlow, T. B. & Davis, R. J. 1994, ApJ, 421, L51

Liedahl, D. A., Osterheld, A. L. & Goldstein, W. H. 1995, ApJL, 438, L115

Owocki, S. P., Castor, J. I. & Rybicki, G. B. 1988, ApJ, 335, 914

Parkin, E. R. & Pittard, J. M. 2010, MNRAS, 406, 2373

Parkin, E. R., Pittard, J. M., Hoare, M. G., Wright, N. J. & Drake, J. J., 2009, MNRAS, 400, 629

Preibisch, T., Balega, Y. Y., Schertl, D. & Weigelt, G. 2002, A&A, 392, 945

Proga, D., Stone, J. M. & Drew, J. E. 1998, MNRAS, 295, 595

Reipurth, B. & Bally, J. 2001, ARA&A, 39, 403

Sim, S. A., Drew, J. E. & Long, K. S. 2005, MNRAS, 363, 615

Terebey, S., Shu, F. H. & Cassen, P. 1984, ApJ, 286, 529

Wang, J., Townsley, L. K., Feigelson, E. D., Getman, K. V., Broos, P. S., Garmire, G. P. & Tsujimoto, M. 2007, ApJS, 168, 100